Reviewer comments and responses

Comments to the Author  
The manuscript describes the analysis of a sample of field stars   
with K and Mg abundances that seem to resemble those observed   
in some extreme globular clusters (like NGC2419 and NGC2808).   
Even if the topic is extremely interesting and an explanation of   
these peculiar stars is still lacking, several points, discussed below,   
should be clarified and better discussed before to consider the paper   
for the publication. I find the description of the analysis   
quite qualitative and not fully convincing. Additional details   
should be provided in order to demonstrate the reliability of the   
results. In particular, I'm not fully convinced that K abundances   
are totally free from bias or systematics.  
I suggest to move the paper to the main journal in order to  
provide all the needed information about the analysis.

**The paper has been lengthened for submission to the main journal to accommodate a more thorough description of the analysis. Additional analysis has been included relating to the treatment of possible systematic effects, and relevant technical details have been included.**  
  
- The abundances from the resonance K lines at 7665 and 7699 A are   
highly sensitive to the adopted microturbulent velocities.  
Also, they can be affected by significant NLTE corrections.   
These two effects are not discussed throughout the paper and they   
could be relevant (in particular to understand whether these   
stars are really K-enriched).   
Which value of microturbulent velocities has been adopted?   
I assume that, considering the low spectral resolution of LAMOST,   
microturbulent velocities have been assumed or derived from   
some parametrization (and not measured from the spectra).  
Are the K abundances corrected for NLTE?  
In principle, the NLTE corrections for K abundances   
are negative (depending on the metallicity and atmospheric parameters),   
pointing out that LTE K abundances are over-estimated.

**We thank the referee for raising this point. Microturbulent velocities were adopted from a parameterisation for giant stars, which we have now explicitly included in the updated manuscript. The quoted K abundances assume LTE in order to make comparisons with other literature sources that only quote LTE abundances. We have updated the manuscript to discuss the level of non-LTE corrections, and included figures of [K/Fe] against stellar parameters to demonstrate that there is no strong systematic trend.**

**We note that our data-driven model encapsulates the typical flux for a star with a given effective temperature, surface gravity, metallicity, et cetera, and we are identifying K-enhanced candidates as stars with significant enhancements away from this model. That is to say that when we identify a star as being enhanced in potassium, we are stating that is showing flux enhancements at the potassium line that are larger than what we would expect for a star of that effective temperature, surface gravity, and metallicity. So while non-LTE effects may over-estimate the abundances we derive, non-LTE effects would not result in false positive detections of K-enhanced stars.**

**Microturbulent velocities used for the K abundances are included, and NLTE coprrections are explicitly dealt with. Essentially, due to the data-driven nature of the model, as opposed to a physical model, no LTE assumption is explicitly made; the NLTE effects are taken into account by the model used to fit the spectra, which takes Teff, logg and metallicity as parameters. Therefore in principle there should be no artificial NLTE effects; this is now demonstrated through showing figures of [K/Fe] against logg and Teff, as suggested.**  
  
- I think that the definition of K-enriched, Mg-depleted stars for   
the discussed targets is not totally correct.   
The K abundances are clearly higher than those observed in other stars   
of similar metallicities, but the [Mg/Fe] abundances are compatible with   
those stars. Also, the comparison between these stars   
and the APOGEE sample (Fig.3) suggests that the targets have   
the same Mg level of the APOGEE stars with similar [Fe/H].   
In other words, they are normal stars in terms of Mg but not in terms of K.  
In principle, I don't see any reason to consider   
these stars depleted in Mg. The Mg-poor stars observed in   
NGC2419 and NGC2808 show a depletion in Mg with respect to   
the Mg abundance observed in most of the cluster stars.   
The LAMOST targets can be consider as Mg-poor stars if they show   
abundances significantly lower than those of the majority of the stars   
with similar metallicity.  
  
**The authors acknowledge that further analysis is required to definitively label the candidate stars as Mg-depleted and K-enhanced. To this end, Mg abundances have been estimated for both the candidate stars and the entire 450000 star sample by applying The Cannon to predict the Mg abundance similarly to the technique Ho et al (2017) applied for the prediction of [alpha/Fe]. The model parameters used for the prediction of [Mg/Fe] are taken to be the same as those used to predict [alpha/Fe], resolving the need to re-train the code; instead, The Canon is applied with [Mg/Fe] as a label over the Mg lines only, with the normal labels elsewhere. Details of this process are included in the revised manuscript.**

**We also note that this effect may be real. No stars have been found to have significantly enhanced [K/Fe] \*or\* depleted [Mg/Fe] outside of NGC 2419. Even the stars in NGC 2808 do not show a very strong signature, only a correlation between the abundance ratios. It is possible that whatever mechanism produced the signature in NGC 2419 is then mixed with Mg-rich supernovae ejecta to make more metal-rich stars with more typical [Mg/Fe] and still enhanced [K/Fe]. However, we avoided making these speculative claims until detailed chemical abundances can be measured for a large sample.**

- The abundances of Fe, K and Mg of the entire sample should be   
shown and not only those of the stars selected as K-rich, Mg-poor.   
It is crucial to show also the other "normal" stars.   
If your analysis procedure is not affected by any bias, most of the stars   
should match with the APOGEE stars in terms of Mg and K.  
The definition of "K-rich stars" should be done with respect to a sample   
of stars analysed in the same way.

**Abundances of [K/Fe] and [Mg/Fe] are estimated using The Cannon for the entire sample. For the case of [Mg/Fe], re-training was not necessary due to the similarity between the behaviour of [alpha/Fe] and [Mg/Fe]. For [K/Fe], the model is re-trained to use [K/Fe] as a label over the range of the spectrum where [K/Fe] lines are present, using APOGEE stars with [K/Fe] already known as the training set.**

- Please show the comparison between two LAMOST spectra with similar   
parameters and metallicities but different K abundances, in order to   
demonstrate a clear difference in the strength of the K lines.

**We have added a figure showing two LAMOST spectra with similar stellar parameters and different Mg/K abundances.**

- The authors should show the behaviour of [Mg/Fe] as a function of [Fe/H]   
(similar to Fig, 4 for K)

**We have added this figure (now Figure X).**

- The authors should show the behaviour of [K/Fe] as a function   
of Teff and logg, in order to demonstrate that there are no spurious   
effects due to the analysis (uncertain parameters, NLTE effects...).

**We have added this figure (now Figure X).**

- It is not totally clear what the authors mean with "amplitude A".   
Is it like an equivalent width? Which are the units of these amplitudes?

**We failed to adequately explain this point. The amplitude is the strength of the Gaussian profile used to fit the flux residuals. It is analogous to the depth of the absorption line, and is related to the equivalent width by EW = 2\pi{}A\sigma, where \sigma is the width of the profile. We have updated the manuscript to more clearly explain the profile amplitude.**

- The abundances of Fe, Mg and K for the three MIKE targets should   
be explicitly provided.

**We have included a table of chemical abundances measured from the Magellan/MIKE spectra.**  
  
- A table with the main information about the K-rich stars (coordinates, magnitudes,   
atmospheric parameters) should be provided.

**We have added a table of properties of the candidates, which includes coordinates, magnitudes, stellar parameters, and other ancillary information.**

- The oscillator strengths for the used K lines should be quoted,   
in order to highlight possible systematics with other works.

**Good catch! We have included these in the updated manuscript.**

- The authors suggest that the mechanism responsible of the K-enhancement   
in these stars should be the same responsible of the K-rich, Mg-poor stars   
in some extreme globulars. However, the K enhancement in NGC2419 and NGC2808   
is explained in the framework of the self-enrichment process occurring   
in the early stage of life of the globulars and related to their high   
stellar density. How is possible that the same mechanism occurs in   
field stars?

**We do not claim to have a complete answer to the Mg-K anti-correlation puzzle, but rather seek to provide information that might guide attempts to explain the phenomena. For example, this work is the first that shows that K-rich (and somewhat Mg-poor) stars exist at metallicities other than [Fe/H] ~ -2. If these stars are related to the mechanism responsible in NGC 2419, then detailed abundances for any of our stars will help guide efforts to explain the puzzle.**

**We discuss possible mechanisms only in the context of existing literature. This has been limited to globular clusters because the abundance signature has never been identified elsewhere. The origin of the abundance signature remains unknown. While it may be that self-enrichment is why there are so many peculiar stars in NGC 2419, it does not necessarily follow that the progenitor star could not form elsewhere and pollute other stars outside a globular cluster environment. It would seem unlikely that this extremely rare signature would have two unrelated origins that depend on whether it is found in a cluster environment or a galactic environment. All of this is speculation, which we have tried to avoid, but we aimed to discuss our results within the context of the existing literature.**

- In Section 3.3 the authors claimed that the behaviour of [K/Fe] with   
the metallicity could be due to some artifacts of the analysis.  
Which artifacts? This sentence seems to suggest possible systematics   
in the analysis that should be better investigated.

**We have expanded the discussion on this point. In short, it is because we measure [K/Fe] from the flux residuals (observed flux - data driven model), but some [K/Fe] contribution may already be accounted for in the data-driven model, because the data-driven model just encapsulates “the typical flux for a star of the given temperature, surface gravity, metallicity, etc”. This potential bias is probably small (far smaller than the uncertainties from measuring abundances from LAMOST data!), but we comment on it for completeness.**

**This claim is elaborated on, and the [K/Fe] abundances derived through spectral synthesis (based around the data-driven model) are compared with those estimated by The Cannon.**

- Concerning the check about [Na/Fe], the authors could investigate   
possible trend of [Na/Fe] with the radial velocity. I expect that   
the stars with Na lines contaminated by the interstellar features   
will be those with radial velocities close to 0 km/s.

**This was our intuition too. KEMP TO UPDATE.**

- In the Introduction, where the authors discuss about K abundance   
in other clusters, please quote also the works by Mucciarelli et al. 2017   
and Cerniauskas et al. 2017 that present abundances of K in samples   
larger than those discussed by Carretta et al. 2013.

**These works are now referenced in the introduction. Thank you for raising this.**  
  
- Caption of Figure 4. The authors claim that the individual metallicities   
of the stars in NGC2808 are not available. The metallicities used by   
Mucciarelli et al. 2015 are those of the same stars analysed by   
Carretta et al. 2009.

**These metallicities are now included in the relevant figure. Thank you for raising this.**

**The authors wish to thank the referee for their time, and the valuable and constructive feedback they have provided. The scientific value and robustness of the paper has been significantly increased as a result.**

**Green=done**